

Cosmological and Particle Physics Constraints on a New Non-Abelian SU(3) Gauge Model for Ordinary/Dark Matter Interaction

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We propose a mirror model for ordinary and dark matter that assumes a new SU(3) gauge group of transformations, as a natural extension of the Standard Model (SM). A close study of big bang nucleosynthesis, baryon asymmetries, cosmic microwave background bounds, galaxy dynamics, together with the Standard Model assumptions, help us to set a limit on the mass and width of the new gauge boson. The cross section for the elastic scattering of a dark proton by an ordinary proton is estimated and compare to the WIMP–nucleon experimental upper bounds. It is observed that all experimental bounds for the various cross sections can be accommodated consistently within the gauge model. We also suggest a way for direct detection of the new gauge boson via one example of a SM forbidden process: $e^+ + p \rightarrow \mu^+ + X$, where $X = \Lambda$ or Λ_c .

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I. INTRODUCTION

The mass density ratios computed from the Wilkinson Microwave Anisotropy Probe (WMAP) [1] data show that the present day dynamics of the Universe is driven essentially by the Dark Energy (DE), see e.g. [2]. Indeed, while $\Omega_{DE} = 0.734 \pm 0.029$ for ordinary baryonic matter, i.e. nuclei and electrons, $\Omega_b = 0.0449 \pm 0.0028$ which is around 5 times smaller than the corresponding value for dark matter (DM) $\Omega_{DM} = 0.222 \pm 0.026$.

The nature of dark matter (DM) is a fundamental problem in modern physics. Dark matter, see e.g. [3–6], is a form of matter that does not interact significantly with ordinary baryonic matter. Experimental evidence for dark matter comes from the anisotropies of CMB and the dynamics of galaxy clusters. Elementary particle theory offer scenarios where new particles such as Weakly Interacting Massive Particles (WIMPs), Sterile Neutrinos, Axions, Supersymmetric Particles, etc. are possible candidates for DM. Experimental searches have set limits on the masses and interactions of some of these hypothetical extra particles [7–21].

A possible scenario for dark matter is the presence of a mirror(s) sector(s) of particles [22–26] where the mirror sectors are copies of the Standard Model (SM). If the mirror sectors are no exact copies of the Standard Model, with, e.g. the mirror particles having different masses and/or couplings than the corresponding SM particles one speaks about asymmetric DM models, see e.g. [37–40]. Anyway, ordinary and mirror particles are weakly coupled. Different mirror models provide different mechanisms for the coupling between ordinary matter and DM.

The current letter describes a mirror model which relies on a symmetry which was so far unexplored. Classifying the fundamental matter fields of the Standard Model according to their electric charge leads, quite naturally, to an $SU(3)$ symmetry, which can be made local to give dynamics to the interaction. The model does not requires an *a priori* number of mirror sectors. However, if the dark sectors are exact copies of the SM, to explain the relative abundance between ordinary and dark matter, five dark sectors are required. Note that using the quoted values for Ω_{DM} and Ω_b it follows that $\Omega_{DM}/\Omega_b = 4.94 \pm 0.66$; the error on the ratio was computed assuming gaussian error propagation. Of course, besides the relative abundance the model should be made compatible with the known cosmological constraints, with Big Bang Nucleosynthesis (BBN) and with the experimental bounds on the cross sections for the interaction with ordinary matter.

The gauge model discussed here assumes a new SU(3) symmetry and introduces a new weakly interacting massive gauge boson (WIMG) which couples the different sectors and, in this way, provides the link between dark and ordinary matter. The WIMG, being a massive boson, leaves unchanged the long distance properties of the Standard Model and gravity. Further, to describe the interaction between ordinary and dark matter, the model requires only two new parameters, namely the new gauge coupling g_M and the WIMG mass M .

In order to provide mass to the WIMG, one cannot rely on the Higgs mechanism which demands the breaking of the underlying gauge symmetry. Instead, we introduce a new mechanism for the generation of mass, which requires a scalar in the adjoint representation with a non-vanishing two point condensate. As discussed below, the expectation value

of the boson condensate is gauge invariant and it turns out that the WIMG mass is proportional to this condensate and, therefore, is also gauge invariant.

The paper is organized as follows. In section II the gauge model is described. The problem of the mass generation and the relation with the SM are explored.

II. A NON-ABELIAN GAUGE MODEL FOR DARK MATTER

At energies much larger than the typical electroweak scale, the Standard Model matter fields behave like massless particles. At such high energy scales, it is natural to group the matter fields according to their electric charge

$$\begin{aligned} Q_1 &= \begin{pmatrix} u \\ c \\ t \end{pmatrix}, & Q_2 &= \begin{pmatrix} d \\ s \\ b \end{pmatrix}, \\ Q_3 &= \begin{pmatrix} e \\ \mu \\ \tau \end{pmatrix}, & Q_4 &= \begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix}. \end{aligned} \quad (1)$$

From the point of view of the electric charge, each multiplet can be rotated to produce a physically equivalent multiplet. This property suggests a new $SU(3)$ group of transformations, where the Q_f 's belong to the fundamental representation, which is broken at the electroweak scale. At the level of the lagrangean density the breaking of the new symmetry shows up in the mass matrix whose origin is linked with the Higgs mechanism within the SM. Although, in the SM the fermion fields are chiral fields, in the definition of the Q chirality we will explore the possibility of having non-chiral fields. Therefore, in the model described below, when exploring the implications for particle physics phenomenology, we will consider two cases in what concerns the chiral properties of Q_i : (i) the matter fields in Q_f do not include a chiral projector in their definition, named non-chiral theory below; (ii) the fields in Q_f are all left-handed, called chiral theory below, and a $\gamma_L = (1 - \gamma_5)/2$ should be attached to each field in (1).

The set of fields Q_f in (1) are the ordinary SM matter. In the following, we will assume that DM has a similar structure as observed for ordinary matter. Each DM sector has 4 multiplets which mimic (1) and each sector has its own copy of the SM. For the moment, the number of DM sectors needs not to be fixed. Given that DM does not seem to couple with the photon, it will also be assumed that each sector is blind to the remaining sectors. Therefore, each sector has its own copy of the SM, with the the corresponding electroweak sectors bosons coupling only within the sector that they are associated with. Note that we are not assuming that each sector is an exact copy of the known ordinary family sector. Indeed, looking at the SM and its family structure what is observed is that the masses differ, not only within each family, but also when one considers different families. In principle, a similar situation can occur here, with the various DM families having different couplings and different masses than those found for the SM.

A dynamics between the various families can be defined if the new global $SU(3)$ symmetry is made local. This requires the introduction of a new gauge boson, the WIMG, coupling all sectors. Note that it will be assumed that the different sectors will be coupled via WIMG exchange and by the gravitational interaction. In our cold Universe, for a sufficient high WIMG mass or a sufficient small coupling constant, the connection between the different sectors is suppressed or essentially vanishing or is provided by gravitational coupling.

The gauge model includes the WIMG field M_μ^a , the matter fields Q_f , where f is a flavor index, and a real scalar field ϕ^a belonging to the adjoint representation of the $SU(3)$ group. The scalar field is required to provide a mass to M_μ^a and, in this way, it ensures that the WIMG interaction is short distance.

The Lagrangian for the gauge theory reads

$$\begin{aligned} \mathcal{L} &= -\frac{1}{4}F_{\mu\nu}^a F^{a\mu\nu} + \sum_f \bar{Q}_f i\gamma^\mu D_\mu Q_f + \\ &+ \frac{1}{2}(D^\mu \phi^a)(D_\mu \phi^a) - V_{oct}(\phi^a \phi^a) \end{aligned} \quad (2)$$

where $D_\mu = \partial_\mu + ig_M T^a M_\mu^a$ is the covariant derivative, T^a stands for the generators of $SU(3)$ group, m_f the current quark mass matrix and V_{oct} is the potential energy associated with ϕ^a . The second term in (2) includes a sum over all families of matter, ordinary and dark matter, and, within each family, over the four multiplets. We have omitted in \mathcal{L} the terms defining the SM for each family and the terms associated with the quantization of the theory.

A. Scalar Fields and WIMG Mass

In the model, we require the WIMG to be a massive gauge boson to leave the long distance properties of the gravitational interaction unchanged. In order to generate a mass to M_μ^a keeping gauge invariance, one has to rely on scalar fields.

The Higgs mechanism leaves a number of components of M_μ^a massless and, to keep the long distance forces unchanged, the Higgs mechanism must be excluded as a way to give mass to the gauge fields.

In [28], the authors propose a mechanism for mass generation via the introduction of a scalar condensate which complies with gauge invariance and provides the same mass for all the components of the gauge field. In the following, we will assume that the WIMG acquires mass through this mechanism. For completeness we will provide the details of the mass generation mechanism.

The kinetic term associated with the scalar field accommodates a mass term for the WIMG field. The gauge field mass term is associated with the operator

$$\frac{1}{2} g_M^2 \phi^c (T^a T^b)_{cd} \phi^d M_\mu^a M^{b\mu}. \quad (3)$$

The scalar field cannot acquire a vacuum expectation value without breaking gauge invariance. However, to generate a mass for the WIMG it is sufficient to assume a non-vanishing boson condensate $\langle \phi^a \phi^b \rangle$. The origin of this condensate can be associated with local fluctuations of the scalar field.

If the dynamics of the scalar field is such that

$$\langle \phi^a \rangle = 0 \quad \text{and} \quad \langle \phi^a \phi^b \rangle = v^2 \delta^{ab}, \quad (4)$$

given that for the adjoint representation $\text{tr}(T^a T^b) = 3 \delta^{ab}$, it follows that the square of the WIMG mass reads

$$M^2 = 3 g_M^2 v^2. \quad (5)$$

Note that v^2 and, therefore, the WIMG mass are gauge invariant. The proof of gauge invariance follows directly from the transformations properties of ϕ^a .

The vacuum expectation values required to give mass to the WIMG (4) minimize the type of potential used in the Higgs mechanism,

$$V(\phi) = \lambda (\phi^a \phi^a - 8v^2)^2, \quad (6)$$

but without breaking the $SU(3)$ symmetry, as would happen in the case of the standard Higgs mechanism. The relations derived above are valid at the tree level and one should investigate how they are changed by the quantum corrections. However, if the theory has a perturbation solution, as is assumed here, at least in lowest order the mass and vacuum expectation values relations should hold.

The WIMG mass (5) is proportional to the effective gauge coupling g_M . One expects to have a small enough g_M . This does not imply necessarily that M is small. The precise value of M depends on the relative values of v and g_M .

B. Other Symmetries of \mathcal{L}

The Lagrangian density \mathcal{L} has several symmetries, both discrete and continuum, besides the local $SU(3)$ gauge invariance. In the present work we are mainly interested in exploring the implications of (2) for DM and will not explore the full structure "hidden" in \mathcal{L} . However, let us comment on some of these symmetries.

The sectors in \mathcal{L} are copies of the ordinary matter. Therefore, it is possible to assign a baryonic and a leptonic number for each sector and, within each family, they are conserved.

The Lagrangian density (2) is invariant under rotations of Q_f . This $SU(4)$ flavor like symmetry per sector mixes leptons and quarks and is broken explicitly by the mass terms. In the case of symmetric DM, where all sectors are copies of the ordinary matter sector, this invariance defines a larger group, i.e. one can rotate the Q 's not only within each family but also between sectors.

In the SM the generation of the mass for the fermion fields occurs, within in the electroweak sector. The diagonalization of the mass matrix requires a rotation of the fields by a unitary transformation over the up and down type of flavors. In (2), such a rotation introduces an unitary flavor matrix associated with the WIMG vertex. For example, in this way, the model can accommodate neutrino mixing.

The implications of the \mathcal{L} symmetries will be addressed in future publications [27].

III. WIMG PROPERTIES AND COSMOLOGICAL CONSTRAINTS

In this section we described how the cosmological constraints can be accommodated within the non-abelian gauge model.

A. Big Bang Nucleosynthesis and Baryon Asymmetries

The gauge model summarized in (2) has new relativistic degrees of freedom that can increase the expansion rate of the early Universe [29] and affect big bang nucleosynthesis (BBN) [30].

After inflation, the temperature for the thermal baths associated with each particle species is not necessarily the same [29]. It depends on the various possible reactions enabling equilibria and on the Universe thermal history. Let us start discuss the simplest possible picture where all the dark sectors have the same temperature, different from the ordinary matter thermal bath, i.e. we are assuming that asymmetric reheating takes place after inflation as in [31–33].

The number of possible new particles contributing to the radiation density during the BBN epoch are constrained by the ${}^4\text{He}$ primordial abundance and the baryon-to-photon ratio $\eta = n_b/n_\gamma$, where n_b is the baryon density and n_γ the photon density in the Universe [34]. For a radiation dominated Universe at very high temperatures, neglecting the particles masses, the energy and entropy densities are given by [35]

$$\rho(T) = \frac{\pi^2}{30} g_*(T) T^4 \quad \text{and} \quad s(T) = \frac{2\pi^2}{45} g_s(T) T^3, \quad (7)$$

where

$$g_*(T) = \sum_B g_B \left(\frac{T_B}{T} \right)^4 + \frac{7}{8} \sum_F g_F \left(\frac{T_F}{T} \right)^4 \quad (8)$$

and

$$g_s(T) = \sum_B g_B \left(\frac{T_B}{T} \right)^3 + \frac{7}{8} \sum_F g_F \left(\frac{T_F}{T} \right)^3, \quad (9)$$

are the effective number of degrees of freedom during nucleosynthesis, $g_{B(F)}$ is the number of degrees of freedom of the boson (fermion) species $B(F)$, $T_{B(F)}$ is the temperature of the thermal bath of species $B(F)$ and T the temperature of the photon thermal bath.

In our case, we consider that the ordinary and dark sectors are decoupled, just after reheating, with different temperatures: T for ordinary matter and T' for the dark sectors. For the dark sectors, the energy $\rho'(T')$ and entropy $s'(T')$ densities are given as in (7) after replacing $g_*(T) \rightarrow g'_*(T')$ and $g_s(T) \rightarrow g'_s(T')$, i.e. the effective number of degrees of freedom in the dark sector, and replacing T by T' . The entropy in each sector is separately conserved during the Universe evolution, which leads that $x = (s'/s)^{1/3}$ is time independent. Assuming the same relativistic particle content for each sector of the modern universe, one has $g_s(T_0) = g'_s(T'_0)$ and it follows that $x = T'/T$.

For a radiation dominated era, the Friedman equation is

$$H(t) = \sqrt{(8\pi/3c^2) G_N \bar{\rho}}, \quad (10)$$

where the total energy density is given by $\bar{\rho} = \rho + N_{DM} \rho'$, where N_{DM} is the number of dark sectors. From the expression for ρ' , it follows

$$H(t) = 1.66 \sqrt{\bar{g}_*(T)} \frac{T^2}{M_{Pl}}, \quad (11)$$

where

$$\bar{g}_*(T) = g_*(T) (1 + N_{DM} a x^4) \quad (12)$$

and M_{Pl} is the Planck mass. The parameter $a = (g'_*/g_*) (g_s/g'_s)^{4/3} \sim 1$, unless T'/T is very small [29]. At the nucleosynthesis temperature scale of about 1 MeV, the relativistic degrees of freedom (photons, electrons, positrons and neutrinos) are in a quasi-equilibrium state and $g_*(T)|_{T=1\text{MeV}} = 10.75$. The extra dark particles change g_* to

$\bar{g}_* = g_* (1 + N_{DM} x^4)$. The bounds due to the relative abundances of the light element isotopes (^4He , ^3He , D and ^7Li) are usually written in terms of the equivalent number of massless neutrinos during nucleosynthesis: $3.46 < N_\nu < 5.2$ [36]. The extra degrees of freedom introduced by the dark sectors lead to $\Delta g_* = \bar{g}_* - g_* = 1.75 \Delta N_\nu < 3.85$, where ΔN_ν is the variation in equivalent number of neutrinos, and $T'/T < 0.78/N_{DM}^{1/4}$ to reconcile the gauge model with the BBN data. If $N_{DM} = 5$, as required by to explain the observed ratio Ω_{DM}/Ω_b , then $T'/T < 0.52$. In conclusion, the asymmetric reheating mechanism leads always to dark universes which are colder than our one universe.

The baryon asymmetry is parameterized by the baryon-to-photon ratio η . The density number of photons n_γ is proportional to T^3 and, therefore, one can write the density number of dark-photons as $n'_\gamma = x^3 n_\gamma$. The ratio of dark-baryons to ordinary-baryons is given by $\beta = \Omega'_B/\Omega_B = x^3 \eta'/\eta$ [32]. The bounds from the BBN on $x = T'/T$ imply that the baryon asymmetry in the dark sector is greater than in the ordinary one. Indeed, using the upper bound $x \sim 0.78/N_{DM}^{1/4}$ and assuming that each sector contributes equally to the Universe's energy density $\beta \sim 1$, we obtain $\eta' \sim 2.1 N_{DM}^{3/4} \eta$. For the special where $N_{DM} = 5$ it follows that $\eta' \sim 7\eta$. Asymmetric Dark Matter models, see e.g. [37–40], give similar results for the baryon asymmetry.

In principle, the presence of mirror baryon dark matter (MBDM) could give some effect on the CMB power spectrum. The reason is that the acoustic oscillations of MBDM could be transmitted to the ordinary baryons. In Ref. [32] this effect was analyzed and their conclusion is that to obtain an observable effect in CMB data it is necessary to have a ratio of temperatures $T'/T \geq 0.6$. This bound combined with the BBN analysis provides a lower bound for the number of dark sectors: $0.35 < N_{DM}$.

B. Galaxy Dynamics and Long Range Interactions

Galaxy dynamics provide further constraints on DM, see e.g. [41, 42]. In the gauge model there is no direct coupling between the photon and its dark brothers. Further, it is assumed that the different sectors behave as the ordinary matter family. It seems natural that the galaxy halos are neutral relative to the $U(1)$'s within each sector. The observed dark matter halos suggest that DM are effectively collisionless and demand an upper bound in the cross section of DM-DM interactions [44, 47–49]. The T'/T bound estimated from BBN complies with such a statement. A typical cross section is given by $\sigma \approx (g^2 T/\Lambda^2)^2$, where g is the interaction coupling constant, T is the temperature and Λ a typical mass scale of the interaction. If the dark sectors are copies of the ordinary matter sector, i.e. g and Λ are of the same order of magnitude, one can write $\sigma'/\sigma = (T'/T)^2$, where σ' (σ) is the cross section for the dark (ordinary) family. The temperature bound from BBN implies that $\sigma'/\sigma < 0.61/\sqrt{N_{DM}}$ and, as long as T'/T is sufficiently small, DM becomes effectively collisionless.

IV. PARTICLE PHYSICS PHENOMENOLOGY

For the first hadronic ordinary family, the WIMG-quark interaction part of the Lagrangian is

$$\begin{aligned} \mathcal{L}_{Mq} = & \frac{g_M}{2} [\bar{u} \gamma^\mu \Gamma c] (M_\mu^1 - iM_\mu^2) + \frac{g_M}{2} [\bar{u} \gamma^\mu \Gamma t] (M_\mu^4 - iM_\mu^5) + \frac{g_M}{2} [\bar{c} \gamma^\mu \Gamma t] (M_\mu^6 - iM_\mu^7) \\ & + \frac{g_M}{2} [\bar{u} \gamma^\mu \Gamma u] \left(M_\mu^3 + \frac{1}{\sqrt{3}} M_\mu^8 \right) + \frac{g_M}{2} [\bar{c} \gamma^\mu \Gamma c] \left(-M_\mu^3 + \frac{1}{\sqrt{3}} M_\mu^8 \right) + \frac{g_M}{2} [\bar{t} \gamma^\mu \Gamma t] \left(-\frac{2}{\sqrt{3}} M_\mu^8 \right) \\ & + h.c. , \end{aligned} \tag{13}$$

where $\Gamma = I$ for the non-chiral theory and $\Gamma = \gamma_L$ for the chiral theory. The remaining families have similar patterns of interactions. The new vertices give rise to flavor changing type of processes within the same family. Given that the WIMG has no electrical charge, it seems that it can give rise to flavor changing neutral processes which are, at most, suppressed by $\sim g_M^2/M^2$. However, the flavor structure of (13) and given that the WIMG propagator is flavor diagonal, the S -matrix element for these processes vanishes. For example, as discussed below, the WIMG vertices give no contributions to the lepton flavor violation processes reported in the particle data book [35]. In this sense, the WIMG interaction is compatible with the GIM (Glashow-Iliopoulos-Maiani) mechanism and the flavor changing neutral currents should remain suppressed at high energies. However, the interaction Lagrangian (13) allows lepton family number violation, see e.g. figures 2 and 4.

A. Lepton Anomalous Magnetic Moment

The new gauge boson provide corrections to the lepton-photon vertex which contribute to the lepton anomalous magnetic moment as depicted in fig.1. The new contributions to $(g - 2)/2$ due to the WIMG are UV-finite and read

$$a_{e,\mu} = \frac{g_M^2}{16\pi^2} \left(\frac{m_{e,\mu}}{M} \right)^2 \left(\frac{5}{3} - \frac{3}{4} \frac{m_\tau}{m_{e,\mu}} \right) \quad (14)$$

and

$$a_\tau = \frac{7g_M^2}{96\pi^2} \left(\frac{m_\tau}{M} \right)^2, \quad (15)$$

for the non-chiral theory and

$$a_l = \frac{5g_M^2}{96\pi^2} \left(\frac{m_l}{M} \right)^2, \quad (16)$$

where $l = (e, \mu, \tau)$, if the particle in the multiplets are left-handed. In (14), (15) and (16) only the leading contributions in m_l^2/M^2 , where m_l is the lepton mass and M the WIMG mass, are taken into account.

The particle data group [35] quotes the following values for the anomalous magnetic moment $a_l = (g - 2)_l/2 = (1159.65218073 \pm 0.00000028) \times 10^{-6}$ for $l = e$, $a_\mu = (11659208.9 \pm 5.4 \pm 3.3) \times 10^{-10}$ and $a_\tau > -0.052$ and < 0.013 . For the muon there is a 3.2σ difference between the experimental value a_μ^{exp} and the standard model prediction a_μ^{SM} which is of $\Delta a_\mu = a_\mu^{exp} - a_\mu^{SM} = 255(63)(49) \times 10^{-11}$.

For the non-chiral theory the WIMG contribution to the lepton anomalous magnetic moment is, for the electron and for the μ , negative due to the τ loop correction to the vertex. Therefore, in the non-chiral theory, the WIMG cannot explain Δa_μ and $a_{e,\mu}$ should be, at most, of the order of the experimental error. This provides the constraints

$$\frac{g_M^2}{M^2} \leq 6.50 \times 10^{-14} \text{ MeV}^{-2} \quad (17)$$

or

$$\frac{g_M^2}{M^2} \leq 8.14 \times 10^{-13} \text{ MeV}^{-2} \quad (18)$$

if one uses a_e or a_μ ; for a_μ the errors reported in [35] were added in quadrature. In the above calculation we used $m_e = 0.511$ MeV and $m_\mu = 105.658$ MeV. These bounds can be rewritten in terms of the WIMG mass as $M \geq g_M \times 3.9$ TeV and $M \geq g_M \times 1.1$ TeV, respectively.

The WIMG contribution to the τ anomalous magnetic moment should be $\sim 1.5 \times 10^{-9}$ or smaller. For the chiral theory, the WIMG contribution to $(g - 2)/2$ should comply with the above results and can be, at most, of the order of the muon anomaly Δa_μ , i.e. $a_\mu \leq 255(63)(49) \times 10^{-11}$, therefore

$$\frac{g_M^2}{M^2} \leq 4.33 \times 10^{-11} \text{ MeV}^{-2} \quad (19)$$

and the WIMG mass should obey $M \geq g_M \times 0.152$ TeV. For the non-chiral theory, the choice of Δa_μ to define the WIMG mass complies with the experimental error for the electron and tau. Indeed, from (16) it follows that the contribution of the new gauge bosons to the electron/tau magnetic moment is $a_l = (m_l^2/m_\mu^2) a_\mu$. These scaling laws give an $a_e = 6.0 \times 10^{-14}$ and $a_\tau = 7.2 \times 10^{-7}$ which are smaller than the experimental error.

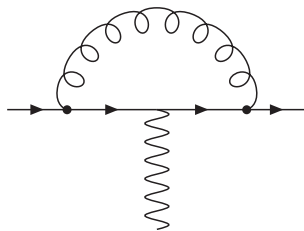


FIG. 1: Lepton-photon vertex correction by WIMG exchange.

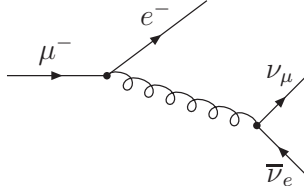


FIG. 2: Muon decay by WIMG exchange process.

B. Lepton Flavor Violation Decays

The Lagrangian (13) allows for flavor changing processes within the same family. The WIMG propagator is flavor diagonal, therefore only those processes where the propagator links the same type of vertices at both ends can have a non-vanishing S -matrix. The following lepton family number violating decays

$$\begin{aligned} \mu^- &\longrightarrow e^- \nu_e \bar{\nu}_\mu, \quad \mu^- \longrightarrow e^- e^+ e^-, \quad \tau^- \longrightarrow e^- e^+ e^-, \\ \tau^- &\longrightarrow e^- \mu^+ \mu^-, \quad \tau^- \longrightarrow \mu^- e^+ e^-, \quad \tau^- \longrightarrow \mu^+ e^+ e^-, \end{aligned}$$

are forbidden within the model because they require different vertices connected by the WIMG propagator. On the other hand, the quark-WIMG vertex structure gives a vanishing S -matrix for $\mu^- \rightarrow e^- \nu_e \bar{\nu}_\mu$. Processes with photons, such as, $\mu^- \rightarrow e^- \gamma$ or $\mu^- \rightarrow e^- \gamma \gamma$ can only occur via loops and are highly suppressed at low energies. The same arguments apply to $B_d \rightarrow e^- \tau^+$ and $B_s^0 \rightarrow \mu^+ \tau^-$ which are forbidden in the model. It turns out that the gauge theory described here complies with the lepton flavor violation bounds reported in the particle data book.

C. Muon Beta Decay

The WIMG can also contribute to the leptonic decays of the μ , the D 's and the B 's mesons. We start by computing the main muonic decay channel $\mu^- \rightarrow e^- \nu_\mu \bar{\nu}_e$ as shown in fig. 2. The S -matrix gets a contribution from W exchange and WIMG exchange. To leading order in M_W and M , for the chiral theory the matrix element for the transition reads

$$\overline{(i\mathcal{M})^2} = 64 G_F^4 \left[1 - \frac{1}{2\sqrt{2}} \frac{g_M^2/M^2}{G_F} \right] (p_\mu \cdot p_{\bar{\nu}_e}) (p_e \cdot p_{\nu_\mu}). \quad (20)$$

For the chiral theory an extra factor of 1/2 should multiply the g_M^2/M^2 . The new contribution modifies the Fermi coupling constant as follows

$$\begin{aligned} G_F &\longrightarrow G_F \left[1 - \frac{1}{2\sqrt{2}} \frac{g_M^2/M^2}{G_F} \right]^{1/4} \\ &\approx G_F \left[1 - \frac{1}{8\sqrt{2}} \frac{g_M^2/M^2}{G_F} \right], \end{aligned} \quad (21)$$

with $G_F = 1.16637(1) \times 10^{-5} \text{ GeV}^{-2}$ [35]. Requiring that the WIMG contribution to be of order of the error on G_F or smaller, then

$$\frac{1}{8\sqrt{2}} \frac{g_M^2}{M^2} \leq 1.0 \times 10^{-10} \text{ GeV}^{-2}$$

or

$$\frac{g_M^2}{M^2} \leq 1.13 \times 10^{-9} \text{ GeV}^{-2}. \quad (22)$$

The relative error on the muon decay width produces a slightly larger g_M^2/M^2 lower limit. The corresponding mass bound is $M \geq g_M \times 28 \text{ TeV}$. Assuming that $g_M \approx e = \sqrt{4\pi\alpha} = 0.30$, we have a lower bound the WIMG mass of $\approx 9 \text{ TeV}$ which complies with both the β -decay, i.e. the error on the Fermi coupling constant, and the muon decay. These mass bounds are more restrictive than the bounds from the anomalous magnetic moment.

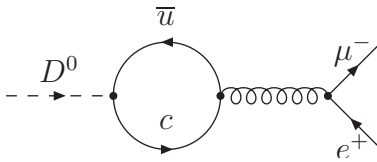


FIG. 3: D^0 decay to (μ^+, e^-) by WIMG exchange process.

D. B and D Leptonic Decays

?

The WIMG vertices can give rise to the following leptonic decays

$$D^0(c\bar{u}) \longrightarrow \mu^- e^+, \quad B^0(b\bar{d}) \longrightarrow \tau^- e^+, \quad B_s^0(s\bar{b}) \longrightarrow \tau^- \mu^+$$

and complex conjugate decays. The widths can be computed using the relation $\langle 0 | \bar{q} \gamma^\mu \gamma_5 q' | (q\bar{q}') \rangle = i f P^\mu$, where $| (q\bar{q}') \rangle$ stands for the meson state composed of quarks $q\bar{q}'$, f is the meson decay constant and P the four-momentum of the meson. Note that the above decays are possible only within the chiral theory. The heavy meson leptonic decay width is calculated by evaluating the amplitude shown in fig. 3, exemplified for $D^0 \rightarrow \mu^- e^+$, which in the general case is given by:

$$\Gamma = \frac{1}{256 \pi} \frac{g_M^4}{M^4} f^2 m_l^2 m_m \left(1 - \frac{m_l^2}{m_m^2} \right)^2, \quad (23)$$

where we have assumed that the lightest lepton is massless, m_l is the mass of the heavier lepton and m_m the mass of the meson state.

For the decay $D^0 \rightarrow \mu^- e^+$, using a $m_{D^0} = 1.864$ GeV and $f_{D^0} = 0.206$ GeV, the bound (22) implies for the corresponding branching ratio

$$Br(D^0 \rightarrow \mu^- e^+) < 8.7 \times 10^{-13} \quad (24)$$

to be compared with the experimental limit [35] of $Br(D^0 \rightarrow \mu^- e^+) < 2.6 \times 10^{-7}$. For the decay $B^0 \rightarrow \tau^- e^+$, using a $m_{B^0} = 5.279$ GeV [35] and $f_{B^0} = 0.22$ GeV [43], the bound (22) gives a

$$Br(B^0 \rightarrow \tau^- e^+) < 2.3 \times 10^{-9} \quad (25)$$

to be compared with the experimental limit [35] of $Br(B^0 \rightarrow \tau^- e^+) < 2.8 \times 10^{-5}$. Finally, for the decay $B_s^0 \rightarrow \tau^- \mu^+$, using a $m_{B_s^0} = 5.366$ GeV [35] and $f_{B_s^0} = 0.24$ GeV [43], the bound (22) gives a

$$Br(B_s^0 \rightarrow \tau^- \mu^+) < 2.7 \times 10^{-9}. \quad (26)$$

Unfortunately, for this decay there is no experimental information. The D^0 and B^0 branching ratios are at least two orders of magnitude smaller than the experimental upper bounds. For B_s^0 , the experimental bounds coming from $g-2$ and muon decay predicts a branching ratio of the same order of magnitude as for B^0 . Note that, in the standard model, the decays discussed are not allowed at tree level but they are one-loop allowed processes. The same type of processes can give rise to the production of dark matter. For example, the bounds (24), (25), (26) also apply to the decays where the leptons are replaced by their dark counter parts. These bounds suggests that the branching rates for production of dark matter from D , B^0 and B_s^0 decays are, at most, of the order of 10^{-9} . Note that the width are proportional to the lepton mass squared and vanish for massless particles in the final state.

V. WIMG PROPERTIES AND WIMG INDUCED PROCESSES

The estimates for the WIMG mass suggest a $M \geq 9$ TeV. Then, from the point of view of the WIMG all the particles in the multiplets (1) are massless. This simplifies considerably the computation of the WIMG width and it follows

$$\Gamma = \frac{g_M^2 M}{24\pi} N_F, \quad (27)$$

where N_F is the number of multiplets. For the chiral theory (27) should be multiplied by 1/2. The bound (22) gives $\Gamma \leq 1.5 \times 10^{-5} M^3 N_F$, where Γ and M are given in TeV. It follows that $\Gamma \approx 1$ TeV or smaller and, therefore, the WIMG should have a very short lifetime $\tau = 1/\Gamma \approx 7 \times 10^{-28}$ s. This means that in the cosmic rays either the WIMG is produced via high energy processes or it is absent from the cosmic rays spectrum.

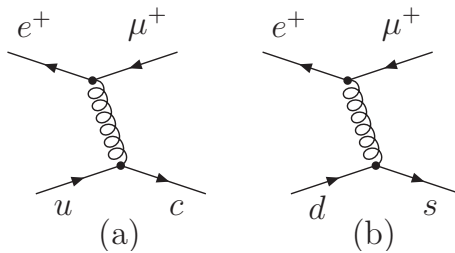


FIG. 4: Lepton-quark scattering with violation of the lepton family number and flavor exchange by WIMG mediated processes. Positron conversion to antimuon and flavor exchange $u \rightarrow c$ (a) and $d \rightarrow s$ (b).

A. A Proposal for WIMG Detection

The WIMG interaction can give rise to processes which are forbidden in the Standard Model. A window to the detection of the new gauge boson is the collision $e^+ + p \rightarrow \mu^+ + X$, where $X = \Lambda$ or Λ_c , which can occur via t -channel WIMG exchange but is forbidden in the Standard Model. At the parton level, the tree-level amplitude for lepton-quark scattering with violation of the lepton family number and flavor exchange is shown in fig. 4 (a) and (b), for charm and strangeness production. The total cross section at the parton level reads

$$\begin{aligned} \sigma(s) &= \frac{1}{384\pi} \frac{g_M^4}{M^4} s \left(1 - \frac{m^2}{s}\right)^2 \left(2 + \frac{m^2}{s}\right) \\ &\leq 8.2 \times 10^{-13} \left(\frac{s}{1 \text{ GeV}^2}\right) \text{ pbarn}, \end{aligned} \quad (28)$$

where s is the c.m energy, m the mass of the quark in the final state and the bound comes from (22).

B. Dark Proton–Ordinary Proton Cross Section

The dark sectors of the gauge model allow for the formation of dark nucleons which can interact with ordinary matter via WIMG exchange. If the dark sectors mimic the ordinary matter sector and given the current temperature of the Universe, the dark protons should be stable particles or, at least, their lifetime should be of the same order of magnitude as the lifetime of an ordinary proton.

The cross section for the elastic scattering of a dark proton by an ordinary proton $\sigma(p_dp)$ is estimated from the one-WIMG exchange process as

$$\sigma(p_dp) \approx \frac{m_N^2}{64\pi} \left(\frac{g_M}{M}\right)^4. \quad (29)$$

For a nucleon mass of $m_N = 1 \text{ GeV}$ and using the bound (22), it follows that $\sigma(p_dp) < 2.5 \times 10^{-48} \text{ cm}^2$.

This upper bound on $\sigma(p_dp)$ can be compared with the experimental bounds for the elastic cross section for the WIMP–nucleon scattering σ_{WIMP} . The current experimental limits are summarized in, for example, Fig. 5 of [17]. For a WIMP mass of the order of a few GeV, the experimental limit gives $\sigma_{WIMP} < 10^{-41} \text{ cm}^2$. The most stringent limit on the cross section occurs for a WIMP mass around 50 GeV, where the corresponding lower limit on the WIMP–nucleon cross section is $7.0 \times 10^{-45} \text{ cm}^2$.

Our estimate of $\sigma(p_dp)$ gives a value which is several orders of magnitude below the experimental limit.

VI. CONCLUSIONS

In the present work a mirror gauge model for the particle interactions between DM and ordinary matter is developed. The model postulates a new SU(3) local symmetry, assumes the existence of various sectors, one being our ordinary matter, which are connected via the exchange of a new gauge boson.

The model is compatible with Big Bang Nucleosynthesis and the recent measurements of the CMB. Further, the dark sectors can be made collisionless if the temperature of the dark sectors is sufficiently lower than the observed

temperature of the visible universe. This difference in the temperature seems to point towards an asymmetric dark matter model, i.e. with the dark sectors not being exact copies of the SM sector.

The model is also compatible with particle physics phenomenology. The well established experimental results on the muon β -decay set a strong constraint on the model parameters. The data clearly provides bounds for g_M/M and set a scale for the WIMG lifetime. Our estimate for the WIMG mass gives a value of the order of the electroweak scale or larger, and a WIMG width of about 1 TeV. Such a large Γ means that the WIMG should not be present in the cosmic rays spectrum unless is produced via the collision of the fundamental particles. Further, we have also checked that the contributions from WIMG exchange to FCNC in the ordinary sector either vanish or are well below the current experimental limits.

The new gauge boson can be associated with processes that are forbidden in the SM. These type of reactions allows for the detection of the new $SU(3)$ symmetry with ordinary matter high energy experiments. As an example of SM forbidden processes we suggest the reaction $e^+ + p \rightarrow \mu^+ + X$ where $X = \Lambda$ or Λ_c .

Finally, the elastic cross for the scattering of an ordinary proton by a dark proton is estimated and compared with the experimental upper bounds for the WIMP–proton cross section [10–19]. The numbers provided combining our estimate with the most demanding bounds coming from the muon beta decay give a $\sigma(p_{dp})$ which is several orders of magnitude lower than the upper bound on the WIMP–proton cross section. Our ordinary and dark matter model allows to compute this type of process involving dark and nuclear particles. In particular, we are currently involved in the calculation of the elastic and inelastic cross sections of dark particles by nucleus which we plan to address in a future work [27].

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